

## KERRBB

Multi-temperature blackbody model for a thin, steady-state, general relativistic accretion disk around a Kerr black hole, developed by Li et al. 2004 (astro-ph/0411583). Includes all relativistic effects and self-irradiation of the disk, and allows for a non-zero torque at the inner boundary of the disk (which is assumed to be at the marginally stable orbit). However, zero torque ( $\eta = 0$ ) is recommended in most cases since it is found that the effect of a nonzero torque can, to a good approximation, be absorbed into a zero torque model by adjusting some of the model parameters. KERRBB is intended as an extension of GRAD, which is limited to a non-rotating black hole and does not include self-irradiation.

**par1** =  $\eta$ , ratio of the disk power produced by a torque at the inner boundary to that arising from accretion ( $0 \leq \eta \leq 1$ ).  $\eta = 0$  corresponds to the case of a standard disk with zero torque at the inner edge

**par2** =  $a$ , dimensionless spin of the black hole ( $-1 \leq a \leq 0.9999$ )

**par3** =  $\vartheta_{\text{obs}}$ , disk inclination angle (deg;  $0 \leq \vartheta_{\text{obs}} \leq 85$ ; 0 for face-on)

**par4** =  $M$ , mass of the black hole (solar units)

**par5** =  $\dot{M}_{\text{eff}}$ , “effective” mass accretion rate ( $10^{18}$  g/sec). When  $\eta = 0$ , it is just the mass accretion rate of the disk. When  $\eta > 0$ , the effective mass accretion rate =  $(1 + \eta)$  times the true mass accretion rate

**par6** =  $D$ , distance (kpc)

**par7** =  $f_{\text{col}}$ , spectral hardening factor,  $T_{\text{col}}/T_{\text{eff}}$ . Should be greater than 1, and is believed to be in the range 1.5 – 1.9 for accretion disks around stellar-mass black holes

**par8** = rflag, a flag to switch on/off the effect of self-irradiation (not allowed to be free). If positive, self-irradiation is included; if  $\leq 0$ , self-irradiation is not included

**par9** = lflag, a flag to switch on/off the effect of limb-darkening (not allowed to be free). If positive, disk emission is limb-darkened; if  $\leq 0$ , disk emission is assumed to be isotropic

**K** = normalization, should be fixed to 1